

Crossing the Phantom divide line in the Chaplygin gas model

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Abstract

The rôle of the interaction in reaching and crossing the phantom divide line in the Chaplygin gas model is discussed. We obtain some necessary properties of the interaction that allow the model to arrive at or cross the phantom divide line. We show that these properties put some conditions on the ratio of dark matter to dark energy density in the present epoch.

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1 Introduction and preliminaries

To describe the present accelerated expansion of the universe [1], many models have been introduced. In one of these pictures nearly 70% of the universe is filled by an exotic smooth energy component with negative pressure dubbed as dark energy. One of the candidates for the dark energy is the Chaplygin gas [2], a perfect fluid whose equation of state (EoS) is given by

$$P_d = -\frac{A}{\rho_d}. \quad (1)$$

A is a positive real constant and P_d and ρ_d are the pressure and the energy density of the dark energy respectively in the rest frame of the fluid. The Chaplygin gas plays a dual rôle: it behaves as a dustlike matter in the early era (i.e. for small scale factor a), and as a cosmological constant at late times (i.e. for large values of a). However, this unified model of dark energy and dark matter, suffers from problems such as production of oscillations or exponential blowup of the matter power spectrum which is inconsistent with observation [3]. Besides this noninteracting model is

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unable to describe the coincidence problem [4], i.e. this model cannot explain why the densities of the dark matter and dark energy are of the same order today. This lies on the fact that energy density of the non interacting cold dark matter (whose EoS parameter is $w_m \simeq 0$) redshifts faster than the energy density of the dark energy sector (whose EoS parameter satisfies $w_d < -1/3$), so the ratio of matter to the dark energy density is expected to decrease rapidly as the universe expands in noninteracting models. In addition this model is inconsistent with the phantom divide line crossing which based on astrophysical data seems to be occurred in the present era [5].

In this paper we assume that the universe is a spatially flat Friedmann-Robertson-Walker space-time which besides the dark energy component, whose EoS is given by (1), is composed of the cold dark matter interacting with the Chaplygin gas (which is taken as the dark energy sector) via the interaction term Q :

$$\begin{aligned} \dot{\rho}_d + 3H\left(-\frac{A}{\rho_d^2} + 1\right)\rho_d &= -Q \\ \dot{\rho}_m + 3H\rho_m &= Q. \end{aligned} \quad (2)$$

ρ_m is the density of the pressureless (cold) dark matter. H is the Hubble parameter and "dot" denotes the derivative with respect to the comoving time. The presence of the interaction source, by permitting the energy exchange between dark matter and dark energy, may alleviate the coincidence problem [6], meanwhile, as we will see, depending on the form of interaction, may allow the phantom divide line crossing.

The Friedmann equations are

$$\begin{aligned} \dot{H} &= -4\pi\left(\frac{-A}{\rho_d} + \rho_d + \rho_m\right) \\ H^2 &= \frac{8\pi}{3}(\rho_d + \rho_m). \end{aligned} \quad (3)$$

The condition $\rho_m + \rho_d = \rho_c := \frac{3H^2}{8\pi}$ guarantees the flatness of the universe. The universe is accelerating if

$$0 < \rho_d < \frac{-\rho_m + \sqrt{\rho_m^2 + 12A}}{2}. \quad (4)$$

In the absence of the interaction, i.e. $Q = 0$, ρ_d is given by

$$\rho_d = \sqrt{A + Ca^{-6}}, \quad (5)$$

where C is a real positive constant. The (EoS) parameter of dark energy is

$$w_d := \frac{P_d}{\rho_d} = -\frac{A}{\rho_d^2}. \quad (6)$$

For (5) we have $-1 < w_d < 0$ for $a \in (0, \infty)$, hence $w_d = -1$ cannot be crossed in this model. As a consequence the EoS parameter of the universe $w = \Omega_d w_d$, where $\Omega_d = \frac{\rho_d}{\rho_c} < 1$, cannot cross the phantom divide line. Using $\lim_{a \rightarrow \infty} \rho_d = \sqrt{A}$, one can show that the ratio of dark matter to dark energy density, defined by

$$r = \frac{\rho_m}{\rho_d}, \quad (7)$$

becomes negligible at late time: $\lim_{a \rightarrow \infty} r = 0$. This was expected because, as was mentioned, cold dark matter redshifts faster than dark energy component. Now to investigate what happens in the presence of the interaction [7], let us consider the simple interaction term

$$Q = 3\Gamma H \rho_d, \quad \Gamma > 0. \quad (8)$$

In this model the solution to the first equation in (2) is

$$\rho_d = \sqrt{\frac{A}{1 + \Gamma} + C a^{-6(\Gamma+1)}}, \quad (9)$$

where C is a real positive constant. In this case in contrast to the noninteracting case where w_d was restricted to the set $(-1, 0)$, in principle the EoS parameter of the Chaplygin gas can cross the phantom divide line ($w_d < -1$), for

$$a > \left(\frac{\Gamma A}{C(1 + \Gamma)} \right)^{\frac{-1}{6(\Gamma+1)}}. \quad (10)$$

Obtaining an exact solution for ρ_m or H is not straightforward, but asymptotic solutions can be derived. Using

$$\frac{dH^2}{dx} = -8\pi \left(-\frac{A}{\rho_d} + \frac{3H^2}{8\pi} \right), \quad (11)$$

where $x := \ln a$, we arrive at $\lim_{a \rightarrow \infty} \rho_c = \sqrt{(\Gamma+1)A}$. In this limit the ratio of dark matter to dark energy, $r := \frac{\rho_m}{\rho_d} = \frac{\rho_c}{\rho_d} - 1$, obeys $\lim_{a \rightarrow \infty} r = \Gamma$. In addition the EoS parameter of the universe satisfies $\lim_{a \rightarrow \infty} w = -1$. In [8], where the rôle of the interaction on dynamical evolution of Chaplygin gas was studied in details, it was shown that this is a stable critical point corresponding to a stable attracting scaling solution in which although the Chaplygin gas crosses the phantom divide line but the universe will enter to a de Sitter phase and the big rip is avoided.

So, the presence of the interaction term, which was first introduced to alleviate the coincidence problem, may allow the dark energy to cross the phantom divide line. In the next section we will find some *necessary* conditions for Q to allow the dark energy to reach or cross the phantom divide line and using an example show that there may be a deep relationship between the coincidence problem and $w_d \simeq -1$ in the present epoch.

2 The rôle of the interaction in crossing the phantom divide line and the coincidence problem

In the presence of an arbitrary interaction, obtaining an exact solution to the Friedmann equations (to see whether the model can reach or cross $w_d = -1$), may not be generally possible. In this part, instead of solving these equations, we try to obtain some *necessary* conditions which must be satisfied if the model reach $w_d = -1$. These conditions put some constraints on Q and the parameters of the model. Besides, depending of the form of Q , w_d may be a decreasing function at $t = t_0$, where t_0 is defined by $w_d(t_0) = -1$. In this situation w_d crosses the phantom divide line.

In the Chaplygin gas model, at $t = t_0$ (where $w_d(t_0) = -1$), using (2) and (6) we obtain

$$\dot{\rho}_d(t_0) = -Q(t_0), \quad \rho_d(t_0) = \sqrt{A}. \quad (12)$$

Note that nowadays (i.e. $t \simeq t_0$) the dominate part of the universe is composed of dark sectors. In addition based on astrophysical data [4] : $\frac{\rho_d(t_0)}{\rho_m(t_0)} = \frac{7}{3}$, therefore the second equation in (2) for a real Hubble parameter requires $\rho_d(t_0) > 0$.

The time derivative of the EoS parameter, (6), at $t = t_0$ reduces to

$$\dot{w}_d(t_0) = -2A\rho_d^{-3}(t_0)Q(t_0) = -2Q(t_0)\rho_d^{-1}(t_0). \quad (13)$$

If $Q(t_0) = 0$ then $\dot{w}_d(t_0) = 0$. In this case, from (2), we deduce $\dot{\rho}_d(t_0) = 0$ and to obtain the first nonzero time derivative of w_d at $t = t_0$ we must continue this procedure: By taking another time derivative from (6), and then using $\ddot{\rho}_d(t_0) = -\dot{Q}(t_0)$, which follows from (2) when $w_d = -1$, $\dot{w}_d = -1$ and $\dot{\rho}_d = 0$, we obtain

$$\ddot{w}_d(t_0) = -2\rho_d^{-1}(t_0)\dot{Q}(t_0). \quad (14)$$

If again $\dot{Q}(t_0) = 0$ we get $\ddot{w}_d(t_0) = 0$. For a differentiable Hubble parameter, this method can be continued to give $\rho_d^{(n)} = -Q^{(n-1)}$, $w_d^{(n)}(t_0) = 2A\rho_d^{-3}(t_0)\rho_d^{(n)}(t_0)$, where $(n-1)$ is the order of the first non zero derivative of w_d at t_0 . Therefore finally we obtain

$$w_d^{(n)}(t_0) = -2\rho_d^{-1}(t_0)Q^{(n-1)}(t_0). \quad (15)$$

If the dark energy sector reaches (or tends to) $w_d = -1$ from the quintessence phase, depending on whether it will enter in phantom phase or not, one of the following predicates must be true (we assume that at t_0 , w_d is differentiable):

i) w_d tends asymptotically to -1 . In this case, following (15), all the time derivatives of w_d vanish at $w_d = -1$, hence Q and all of his time derivatives must also vanish in this limit:

$$Q(t_0) = 0, \quad \frac{d^n Q}{dt^n}(t_0) =: Q^{(n)}(t_0) = 0, \quad \forall n \in N. \quad (16)$$

For a derivable continuous w_d this can only occur at $t_0 \rightarrow \infty$. E.g. the noninteracting Chaplygin gas model, $Q = 0$, belongs to this category and as we have seen $w_d = -1$ occurs asymptotically. As another example consider

$$Q = \lambda(P_d + \rho_d) = \lambda(w_d + 1)\rho_d, \quad (17)$$

where λ is a constant. This is the interaction considered for some scalar field ($\phi(t)$) models of dark energy or inflaton, giving rise to the interaction $\lambda\dot{\phi}^2$ [9]. It is clear that at $w_d = -1$, we have $Q = 0$. Using (2) one can show that the higher derivatives of Q also vanish at $w_d = -1$, therefore $w_d = -1$ can only occur asymptotically.

ii) The first nonzero derivative of w_d is negative and of odd order at $w_d = -1$, in this case after reaching, w_d crosses $w_d = -1$ line and the Chaplygin gas enters the phantom phase. In this situation, $Q(t_0) > 0$ or, if $Q(t_0) = 0$ at t_0 the first nonzero derivative of Q is positive and of even order

$$Q(t_0) > 0, \quad \text{or} \{Q(t_0) = 0, \quad \text{and} \quad Q^{(2n)}(t_0) > 0, \quad Q^{(k<2n)}(t_0) = 0\}. \quad (18)$$

The example proposed in the previous section (8), belongs to this category. The condition $\Gamma > 0$ guarantees $Q > 0$.

iii) $\dot{w}_d = 0$ at $w_d = -1$, and the first non-vanishing derivative of w_d is positive and of even order . In this case the dark component does not enter in the phantom phase and $w_d = -1$ is the global minimum of w_d . In this situation, with the same method used in ii), it is easy to show that at t_0 we have $Q = 0$ and the first nonzero derivative of Q is negative and of odd order

$$Q(t_0) = 0, \quad Q^{(2n+1)}(t_0) < 0, \quad Q^{(k<2n+1)}(t_0) = 0. \quad (19)$$

In this case the phantom phase is neither accessible for the dark energy component nor for the universe: $w > -1$.

Therefore in the context of the Chaplygin gas model, reaching or crossing $w_d = -1$ may not possible for an arbitrary Q , e. g. following (i) and (ii) and (iii) models with $Q < 0$ do not arrive at $w_d = -1$. Besides, via considering Q as a function of ρ_m and ρ_d , the above conditions give some relationship between the density of dark matter to density of dark energy at t_0 . In this view one can determine or put some conditions on the value of $r = \frac{\rho_m}{\rho_d}$ at the present epoch where based on some astrophysical data $w_d \approx -1$ is happened. So there may be a deep relationship between coincidence problem and $w_d = -1$ crossing observed in the present era.

To illustrate this point, let us consider the interaction [10]

$$Q = H(\lambda_m \rho_m + \lambda_d \rho_d), \quad (20)$$

and study the conditions required to reach $w_d = -1$. Following (i) and (ii) and (iii) we must have $Q \geq 0$ at t_0 . For an expanding universe this implies $\lambda_m r(t_0) \geq -\lambda_d$.

For

$$r(t_0) = -\frac{\lambda_d}{\lambda_m}, \quad (21)$$

$Q = 0$ and $\dot{w}_d = 0$, but using $\dot{\rho}_d(t_0) = 0$ and $\dot{\rho}_m(t_0) = -3H\rho_m(t_0)$ we obtain $\ddot{w}_d(t_0) = -3\lambda_d H^2(t_0)$ and following (ii) (in order that reaching $w_d = -1$ be allowed) we must have $\lambda_d < 0$. In this situation the dark energy component, remains in the quintessence phase. This possibility is excluded when one considers the interaction $Q = \lambda H(\rho_m + \rho_d)$ which is adopted in some papers [11].

If $\lambda_m r > -\lambda_d$ at t_0 , then w_{d0} cross the phantom divide line. In this case for $\lambda_m > 0$, and $\lambda_d < 0$ we obtain a lower bound for $r(t_0)$:

$$r(t_0) > -\frac{\lambda_d}{\lambda_m} \quad (22)$$

at transition time.

Although obtaining an exact solution to the Friedmann equations in the presence of interaction (20) does not seem to be possible but let us examine whether these equations admit the series solution

$$\begin{aligned} w_d &= -1 + w_{d0}t^\alpha + \mathcal{O}(t^{\alpha+1}), \\ \rho_d &= \rho_{d0} + \rho_{d1}t^\beta + \mathcal{O}(t^{\beta+1}), \\ \rho_m &= \rho_{m0} + \rho_{m1}t^\gamma + \mathcal{O}(t^{\gamma+1}), \\ H &= H_0 + H_1t + H_2t^2 + \mathcal{O}(t^3), \end{aligned} \quad (23)$$

at $t = 0$ (we have taken $t_0 = 0$). (1) leads to $\alpha = \beta$, $\rho_{d0} = \sqrt{A}$ and $\rho_{d1} = \frac{w_{d0}\sqrt{A}}{2}$. We have also assumed that $\dot{H}(t_0) \neq 0$ which lies on the assumption $\Omega_d(t_0) \neq 0$. We first consider the solution with $\alpha \geq 2$. Putting (23) into the first equation in (2) and equating the terms with the same power of t , after some calculation we obtain $\lambda_m \rho_{m0} + \lambda_d \rho_{d0} = 0$, and $\alpha - 1 = \gamma$ and subsequently by using the second equation in (2) we obtain $\alpha = 2$ (viz. $\alpha > 2$ is not allowed), $\gamma = 1$, $\rho_{m1} + 3H_0 \rho_{m0} = 0$ and $2\rho_{d1} = -\lambda_m \rho_{m1} H_0$. By collecting these together, the series solution for $\alpha \geq 2$ at $t = 0$ (which is

also consistent with (3)) is obtained as

$$\begin{aligned} w_d &= -1 - 3\lambda_d H_0^2 t^2 + \mathcal{O}(t^3) \\ \rho_d &= \sqrt{A} - \frac{3}{2}\lambda_d H_0^2 \sqrt{A} t^2 + \mathcal{O}(t^3) \\ \rho_m &= -\frac{\lambda_d}{\lambda_m} \sqrt{A} (1 - 3H_0 t) + \mathcal{O}(t^2) \\ H &= H_0 + 4\pi \frac{\lambda_d}{\lambda_m} \sqrt{A} t + \mathcal{O}(t^2), \end{aligned} \quad (24)$$

where the Hubble parameter at $t = 0$ is determined to be $H_0 = \frac{8\pi}{3} \left(1 - \frac{\lambda_d}{\lambda_m}\right) \sqrt{A}$. This solution corresponds to the system characterized by (21).

Note that $\rho_{m1} + 3H_0\rho_{m0} = 0$ in this example leads to $Q(t_0) = 0$. This lies on the fact that $\alpha = 2$ in (23), and therefore the Chaplygin gas remains in quintessence phase: $\dot{w}_d(t_0) = 0$. Hence (13) results in $Q(t_0) = \dot{w}_d(t_0) = 0$. As we will show this is not true for other possible solutions.

For $\alpha = 1$ (corresponding to the aforementioned case $\lambda_m r(t_0) > -\lambda_d$), one can similarly show that the following solution exists:

$$\begin{aligned} w_d &= -1 + \frac{2}{\sqrt{A}} \left((\lambda_m - \lambda_d) \sqrt{A} - \frac{3H_0^2}{8\pi} \lambda_m \right) H_0 t + \mathcal{O}(t^2) \\ \rho_d &= \sqrt{A} + \left((\lambda_m - \lambda_d) \sqrt{A} - \frac{3H_0^2}{8\pi} \lambda_m \right) H_0 t + \mathcal{O}(t^2) \\ \rho_m &= \left(\frac{3H_0^2}{8\pi} - \sqrt{A} \right) + \left((\lambda_m - \lambda_d + 3) \sqrt{A} - \frac{3H_0^2}{8\pi} (\lambda_m + 3) \right) H_0 t + \mathcal{O}(t^2) \\ H &= H_0 + \left(-\frac{3H_0^2}{2} + 4\pi \sqrt{A} \right) t + \mathcal{O}(t^2). \end{aligned} \quad (25)$$

According to the condition (ii), (18) and subsequently (25) are valid when

$$\lambda_m H_0^2 > \frac{8\pi}{3} (\lambda_m - \lambda_d) \sqrt{A}. \quad (26)$$

This solution crosses $w_d = -1$. If $\lambda_m = 0$ we must have $\lambda_d > 0$ which is the same result discussed in the previous section. If $\lambda_d = 0$, only models with $\lambda_m > 0$ are allowed. If we adopt $\lambda_m = \lambda_d = \lambda$, then $w_d = -1$ crossing is possible whenever $\lambda > 0$.

After the dark energy component of the universe passes $w_d = -1$, the EoS parameter of the universe may cross $w = -1$ too. Note that based on the values estimated for w_d and Ω_d , it is also possible that the universe has an EoS parameter satisfying $w = w_d \Omega_d \simeq -1$ even in the present era, e.g. if we adopt $\Omega_d = 0.73$ and $w_d \simeq -1.33$ [12], we obtain $w = -0.97$.

For a general dark energy model it can be shown that

$$w = -\frac{1}{3H} \frac{\dot{\Omega}_d}{1 - \Omega_d} - \frac{Q}{3H\rho_c(1 - \Omega_d)}. \quad (27)$$

In the Chaplygin gas, from $w_d\Omega_d = w$ we derive

$$w = -A \left(\frac{8\pi}{3} \right)^2 \frac{1}{H^4 \Omega_d}. \quad (28)$$

(27) together with (28) form a system (which is not analytically solvable) describing the behavior of w and Ω_d in terms of the scale factor. If the system arrives at $w = -1$ at $t = \tau$ from (28) we obtain

$$\dot{w}(\tau) = \frac{\dot{\Omega}_d(\tau)}{\Omega_d(\tau)}. \quad (29)$$

The universe enters in the phantom phase provided that $\dot{w}(\tau) \leq 0$. By considering (27), this implies

$$Q \geq 3H\rho_c(1 - \Omega_d), \quad (30)$$

at $t = \tau$. In terms of $\Omega_m = \frac{\rho_m}{\rho_c}$ this inequality becomes

$$Q \geq 3H\rho_m \quad (31)$$

at $t = \tau$. Putting (20) in this inequality gives

$$(\lambda_m - 3)r \geq -\lambda_d, \quad (32)$$

which for $\lambda_m > 3$ and $\lambda_d < 0$, gives a lower bound for r at transition time:

$$r \geq -\frac{\lambda_d}{\lambda_m - 3}, \quad (33)$$

while for $\lambda_m < 3$ and $\lambda_d > 0$, $-\frac{\lambda_d}{\lambda_m - 3}$ becomes an upper bound for r . For $\lambda = \lambda_d = \lambda_m$, the upper bound becomes $r < \frac{\lambda}{3 - \lambda}$.

To study the behavior of the system near $w = -1$ as before we may consider the following series solutions for H and Ω_d at $t = \tau$ (note that $\dot{H}(\tau) = 0$),

$$\begin{aligned} H &= h_0 + h_1(t - \tau)^\nu + \mathcal{O}((t - \tau)^{\nu+1}), \quad \nu \geq 2, \quad h_1 > 0 \\ \Omega_d &= u_0 + u_1(t - \tau)^\mu + \mathcal{O}((t - \tau)^{\mu+1}). \end{aligned} \quad (34)$$

μ and ν are the orders of the first nonzero derivatives of H and Ω_d at $t = \tau$ respectively. As we are studying the transition from quintessence to phantom phase we have taken $h_1 > 0$. Using (27) and the equation

$$w = -1 - \frac{2}{3} \frac{\dot{H}}{H^2}, \quad (35)$$

we find that if $\mu \neq 1$ then we must have $\mu = \nu$ (for details see [13]), but inserting the series solution in (27) yields $\nu = \mu + 1$ which results in $\nu = 2$,

and $\mu = 1$. Note that $\nu = 2$ implies that the system cannot remain in quintessence phase. We also obtain $u_1 = (3 - \lambda_m)h_0 + (\lambda_m - \lambda_d - 3)h_0u_0$. Using (29), (2), and (3) we arrive at

$$\begin{aligned} H &= h_0 - \frac{3}{4}h_0^3 \left(\frac{9}{64\pi^2 A} (3 - \lambda_m)h_0^4 + (\lambda_m - \lambda_d - 3) \right) (t - \tau)^2 \\ &\quad + \mathcal{O}((t - \tau)^3) \\ \Omega_d &= \frac{64\pi^2 A}{9h_0^4} + \left((3 - \lambda_m)h_0 + \frac{64\pi^2}{9h_0^3} (\lambda_m - \lambda_d - 3) \right) (t - \tau) \\ &\quad + \mathcal{O}((t - \tau)^2). \end{aligned} \quad (36)$$

h_0 is not determined, but $h_1 > 0$ gives

$$(3 - \lambda_m)h_0^4 < \frac{64\pi^2 A}{9} (3 - \lambda_m + \lambda_d). \quad (37)$$

This inequality, which using $r = \frac{1 - \Omega_d}{\Omega_d}$ can be shown to be the same as (32), determines the relation between the energy density at transition time and the parameters of the model.

3 Conclusion

We considered the interacting Chaplygin gas as the dark energy component of the universe. Some necessary conditions on the interaction, allowing the Chaplygin gas to reach the phantom divide line ($w_d = -1$), were discussed. Using these condition we found that the ratio of dark matter to dark energy density, r , must satisfy some inequalities (or equalities) which can alleviate the coincidence problem, i.e. at $w_d = -1$, r is whether determined in terms of the parameters of the system or a nonzero lower bound can be found for it. Via some examples, we showed that how the form of interaction determines some of the behavior of the system such as crossing the phantom divide line or remaining in the quintessence phase. We also obtained a series solution to Friedmann equations at phantom divide line. At the end via an example we studied the possibility that the EoS of the universe crosses $w = -1$.

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